

DISKS AROUND THE YOUNG STARS TW HYDRAE AND HD 100546 IMAGED AT 3.4 MILLIMETERS WITH THE AUSTRALIA TELESCOPE COMPACT ARRAY

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Received 2003 February 7; accepted 2003 June 18

ABSTRACT

We present observations of the young stars TW Hya and HD 100546 made with the Australia Telescope Compact Array at 89 GHz with $\sim 2''$ resolution and ~ 3 mJy continuum sensitivity. Compact thermal dust continuum emission is detected from disks surrounding both stars. HD 100546 also shows hints of extended emission, presumably a residual protostellar envelope, which is also visible in scattered light at optical wavelengths. For TW Hya, $\text{HCO}^+ J = 1-0$ line emission from the circumstellar disk is detected and spatially resolved. The observed size and intensity are in good agreement with model calculations based on an irradiated disk with substantial depletions derived previously from single-dish observations of higher J HCO^+ transitions.

Subject headings: circumstellar matter — planetary systems: protoplanetary disks — stars: individual (TW Hydrae, HD 100546)

1. INTRODUCTION

Many young stars exhibit emission from circumstellar dust particles distributed in disks with properties similar to the early solar system, and much activity is currently devoted to characterizing the physical properties of these disks to extract information on planet formation. Observations at millimeter wavelengths are especially important because the disk material beyond a few stellar radii is at temperatures from a few hundreds to a few tens of degrees, and the physical and chemical conditions can be probed in detail in this part of the spectrum (see reviews by Beckwith & Sargent 1996; Wilner & Lay 2000; Langer et al. 2000).

The recent upgrade of the Australia Telescope Compact Array (ATCA) with receivers for the 3 mm atmospheric window provides a new opportunity for high-resolution imaging of protoplanetary disks located in the southern sky that are difficult or impossible to observe with the millimeter arrays located in the northern hemisphere. The ATCA millimeter-wave upgrade is in progress, and the completed facility will ultimately rival the northern interferometers in sensitivity and angular resolution (see Wong & Melatos 2002 for a more complete description). Now, three of the array antennas are equipped with 3 mm band InP MMIC receivers, and an interim local oscillator system allows for tuning in two ranges between 85 and 91 GHz. We present interferometric observations of dust continuum and $\text{HCO}^+ J = 1-0$ line emission at 89 GHz from two southern disk targets: (1) TW Hya, the closest known classical T Tauri star, and (2) HD 100546, a nearby Herbig Be star whose infrared spectrum shows crystalline silicates, indicative of comet-like dust.

TW Hya has received a lot of recent observational attention, largely on account of its close proximity. At a distance of 56 ± 7 pc (*Hipparcos*), TW Hya is almost 3 times closer than the classical T Tauri stars associated with nearby dark clouds like Taurus and Ophiuchus, which makes it an especially attractive target for observations with high angular resolution. At an age of ~ 10 Myr, the TW Hya disk appears to be substantially evolved, with indications of significant grain growth. TW Hya is isolated from any molecular cloud but retains a face-on molecular disk visible in scattered light (Krist et al. 2000; Trilling et al. 2001; Weinberger et al. 2002) that extends to a radius of at least $3''5$ (200 AU). The inner disk of TW Hya has been resolved in dust emission by the VLA at 7 mm (Wilner et al. 2000), and its spectral energy distribution has been modeled extensively (Calvet et al. 2002). The available data are well fitted by an irradiated accretion disk with a developing gap of radius ~ 4 AU, inside which dust coagulation and settling have rendered the disk optically thin.

The isolation of the TW Hya system from any surrounding molecular cloud material facilitates the study of its physical and chemical structure. The TW Hya disk has been detected in a suite of molecular species (CO , HCN , CN , and HCO^+) at submillimeter wavelengths using single-dish telescopes (Kastner et al. 1997; van Zadelhoff et al. 2001). More recent, deeper observations have also detected emission from H^{13}CO^+ and DCO^+ , which indicate deuterium fractionation in the disk similar to cold cores and pristine cometary material (van Dishoeck, Thi, & van Zadelhoff 2003). So far, no observations have been made that spatially resolve the molecular gas surrounding TW Hya. Such spatially resolved observations are important to verify

underlying constructs of the models used to interpret the single-dish data, since the physical conditions are subject to strong radial and vertical gradients, which result in considerable chemical complexity.

HD 100546, among the nearest Herbig Ae/Be stars at 103 ± 6 pc (*Hipparcos*), appears near the dark cloud DC 296.2–7.9 (Hu, The, & de Winter 1989) and shows a disk-like scattered light distribution (Pantin, Waelkens, & Lagage 2000; Augereau et al. 2001; Grady et al. 2001) of substantial size, about $8''$ (8000 AU), as well as strong millimeter emission from dust (Henning et al. 1994, 1998). This relatively isolated system is thought to have an age of ~ 10 Myr (van den Ancker et al. 1997). However, analysis of the full spectral energy distribution suggests the presence of an extended envelope (Henning et al. 1994). Mid-infrared spectroscopy from the *Infrared Space Observatory (ISO)* shows remarkably strong crystalline silicate bands, similar to those observed in comet Hale-Bopp (Malfait et al. 1998) and indicative of unusual and substantial processing of the dust within the disk (Bouwman et al. 2003). Millimeter interferometry has the potential to show directly the presence of a disk component that contains the bulk of the system mass.

2. OBSERVATIONS

We observed TW Hya and HD 100456 at 89 GHz (3.4 mm) with the ATCA during the 2002 austral winter using two compact configurations of three 22 m diameter antennas equipped with millimeter receivers. Table 1 summarizes the observational parameters.

The ATCA observations provided six independent east-west baselines with lengths ranging from 22 to 230 m, resulting in $\sim 2''$ resolution. For each observation, the target pointing center was set to be $5''$ west of the star position. The digital correlator was configured for two dual polarization bands: (1) a wide band with 33 channels spanning the maximum bandwidth of 128 MHz for continuum sensitivity, and (2) a narrow band with 256 channels over 16 MHz to provide high-frequency resolution on the $\text{HCO}^+ J = 1-0$ line, which is one of the few potentially detectable spectral

lines accessible to the interim system. The sky signal is mixed with a Gunn oscillator locked at 80505.5 MHz and passed through an X-band (8000–10800 MHz) filter and splitter module, which results in an accessible frequency range of 88.506–91.305 GHz. A new local oscillator system under construction will allow coverage of the frequency range 85–105 GHz, with a possible extension to 115 GHz.

All calibration and imaging was performed with the MIRIAD software. Complex gains were derived from frequent observations of nearby quasars. Because the narrow band provided low signal-to-noise ratios on the quasars, a phase offset between the wide band and narrow band was determined from short observations of a strong source on each day, either J0423–012 or J0538–440. These short observations of strong sources were also used to determine the bandpass response. For TW Hya, additional weaker quasars located close to the star were included in the observing sequence to provide an empirical check on the atmospheric seeing and effectiveness of the phase calibration. Typical system temperatures were 300–400 K (SSB), with higher values at low elevations. The flux densities were set with reference to the planet Mars; the scatter in derived fluxes on consecutive days suggests uncertainties of approximately 20%.

Images were made of the sum of the two linear polarizations using natural weighting to obtain best sensitivity. All images were cleaned to a cutoff of twice the rms noise level. Because Doppler tracking was not applied during the observations, the correspondence between individual frequency channels and velocity changes with time during the course of the observations. Therefore, the spectral data were imaged in resampled velocity bins of 0.5 km s^{-1} width, coarse compared to these changes, rather than individual frequency channels.

3. RESULTS

3.1. TW Hya

Figure 1 shows the 89 GHz continuum emission detected from TW Hya. The position is coincident with the optical

TABLE 1
OBSERVATIONAL PARAMETERS

Parameter	TW Hya	HD 100546
Observation dates	2002 Jun 1, Aug 7	2002 May 31, Aug 8
Minimum/maximum baseline (m)	22–230	22–230
Pointing center (J2000) α	11 ^h 01 ^m 52 ^s .31	11 ^h 33 ^m 26 ^s .40
Pointing center (J2000) δ	$-34^\circ 42' 17.0'' 3$	$-70^\circ 11' 41'' 2$
Phase calibrator	J1147–381	J1058–800
Flux calibrator	Mars	Mars
Primary beam HPBW	$37''$	$37''$
Synthesized beam HPBW	$6''.9 \times 2''.1$, P.A. 18°	$3''.2 \times 2''.2$, P.A. 16°
K/Jy	10.6	21.1
rms (continuum image) (mJy beam^{-1})	3	3
Detected flux ^a (mJy)	41 ± 4	36 ± 3
Spectral line correlator	257 channels, 16 MHz	
Species/transition	$\text{HCO}^+ J = 1-0$	
Frequency	89.188518 GHz	
Center velocity	0 km s^{-1}	
Channel spacing	0.21 km s^{-1}	
rms (line images) (mJy beam^{-1})	60	60

^a Includes random error only; the systematic uncertainty in the flux scale is estimated to be $\sim 20\%$.

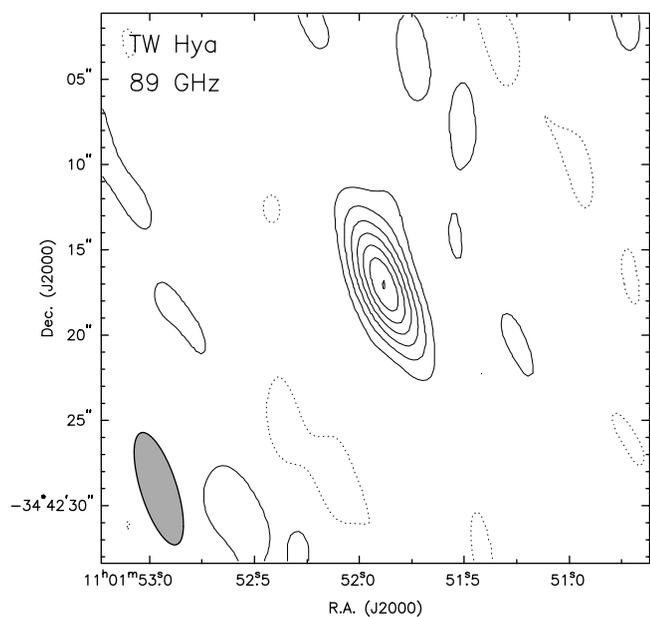


FIG. 1.—TW Hya continuum emission at 89 GHz. The contour levels are $\pm 2, 4, 6, \dots \times 2.5$ mJy. Negative contours are dotted. The ellipse in the lower left corner shows the $6''.9 \times 2''.1$ P.A. 18° synthesized beam.

star and radio detections at longer wavelengths. A Gaussian fit to the visibilities gives a flux of 41 ± 4 mJy (random error only) and an apparent size consistent with an unresolved source slightly broadened by the phase noise on the longer baselines. This Gaussian fit is also compatible with the extended disk seen in scattered light, since the dust emission from the circumstellar disk is strongly centrally peaked by the combination of low opacity and increasing column density and temperature toward the star (see Mundy et al. 1996; Wilner et al. 2000).

The TW Hya 89 GHz continuum flux measurement agrees well with expectations from previously reported measurements at both higher (350 GHz) and lower (43 GHz) radio frequencies (Weintraub, Sandell, & Duncan 1989; Wilner et al. 2000; Wilner 2001). In addition, it is consistent with predictions of various disk models based on the radio observations that contain a population of large dust grains to account for the shallow spectral slope with power-law index $\lesssim 3$ observed in this frequency range (Trilling et al. 2001; Calvet et al. 2002). The good fit of the new 89 GHz measurement with expectations gives confidence in the reliability of the ATCA system and the accuracy of the absolute flux scale.

Figure 2 shows a series of images in four velocity bins for the HCO^+ $J = 1-0$ line emission near LSR velocity ~ 3 km s^{-1} . A narrow line emission feature is clearly visible at the stellar position with a velocity and line width commensurate with previously reported molecular line detections from the face-on disk, which have $V_{\text{LSR}} = 2.9$ km s^{-1} and $\Delta V \approx 0.6$ km s^{-1} . Figure 3 shows the spectrum at the continuum position. The peak line flux at 0.5 km s^{-1} resolution is 0.40 Jy, which corresponds to a brightness temperature of 4.3 K in the $6''.9 \times 2''.1$ beam. The line emission from the disk is clearly spatially resolved; a circular Gaussian fit to the visibilities in a 0.5 km s^{-1} bin gives a FWHM size of $3''.2 \pm 0''.8$.

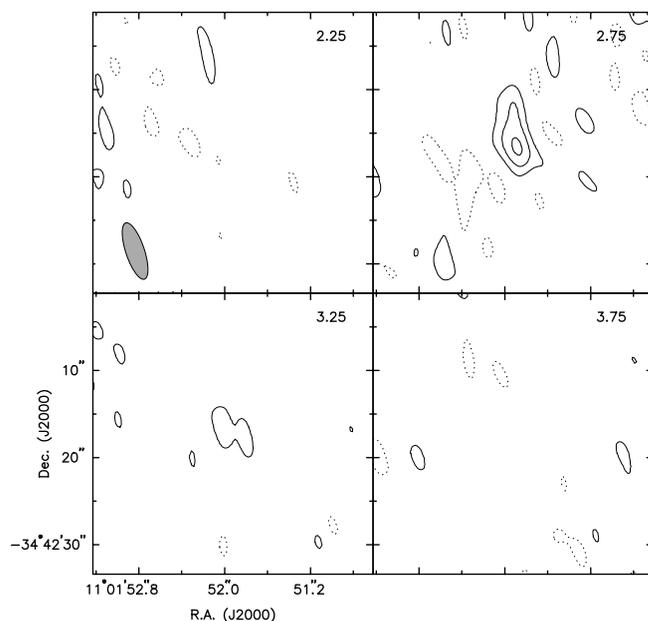


FIG. 2.—Velocity channel maps (width 0.5 km s^{-1}) of HCO^+ $J = 1-0$ line emission observed from TW Hya. The contour levels are $\pm 2, 4, 6, \dots \times 66$ mJy (0.7 K). Negative contours are dotted. The ellipse in the lower left corner of the upper left panel shows the $6''.9 \times 2''.1$ P.A. 18° synthesized beam.

3.2. HD 100546

Figure 4 shows the 89 GHz continuum emission detected from HD 100546. The peak position is coincident with the star to better than $1''$. A fit to the visibilities gives a point source flux of 36 ± 3 mJy at the star position. The residuals to this fit provide a hint that additional emission is present. The image shows an extra peak of low significance located $\sim 7''$ to the southeast. This additional peak may be related to the slight extension seen in the single-dish 1.3 mm bolometer image of Henning et al. (1998). Extensions along the same position angle are also visible in scattered light, in particular in the sensitive *Hubble Space Telescope* STIS images of Grady et al. (2001), where nebulosity can be seen to extend from the northwest to the southeast. No line emission is detected to a limit of 0.17 Jy per channel (3σ), either at the LSR velocity of the cloud DC 296.2–7.9 of 3.6 km s^{-1}

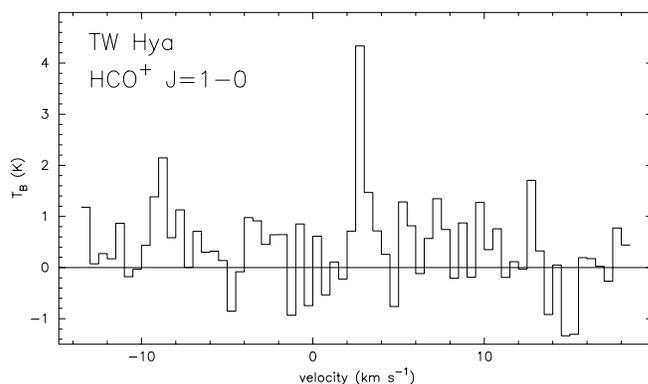


FIG. 3.—Spectrum of HCO^+ $J = 1-0$ line emission observed from TW Hya, at the position of peak continuum emission, binned to 0.5 km s^{-1} velocity resolution, in the $6''.9 \times 2''.1$ P.A. 18° synthesized beam.

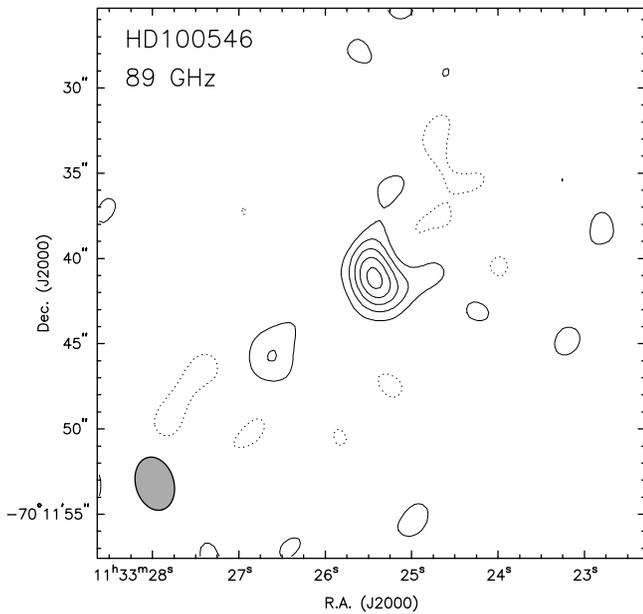


FIG. 4.—HD 100546 continuum emission at 89 GHz. The contour levels are $\pm 2, 4, 6, \dots \times 3.1$ mJy. Negative contours are dotted. The ellipse in the lower left corner shows the $3''.2 \times 2''.2$ P.A. 16° synthesized beam.

(Otrupcek, Hartley, & Wang 2000) or at the stellar LSR velocity of about -1 km s $^{-1}$, which derives from the heliocentric “astrometric” radial velocity of 9 ± 1 km s $^{-1}$ determined from *Hipparcos* (Madsen 2002).

4. DISCUSSION

4.1. TW Hya

The spatially resolved observations of the $\text{HCO}^+ J = 1-0$ emission from TW Hya allow for a check of the basic paradigm that trace molecules are present in an extended upper layer of the disk, with overall abundances depleted with respect to molecular hydrogen by large factors compared to those of interstellar dark clouds (Aikawa & Herbst 1999; Willacy & Langer 2000; Aikawa et al. 2002). In the disk mid-plane, which contains most of the mass, the combination of high densities and low temperatures results in nearly all trace molecules sticking to grain surfaces and disappearing from the gas. At the disk surface, the molecules are photodissociated. In between, below the surface, is a layer that is warm and shielded from stellar radiation and activity, where molecules survive and abundances peak.

The TW Hya disk structure and chemistry has been considered in detail by van Zadelhoff et al. (2001), whose model calculations suggest that the depletion of species like CO and HCO^+ results from a combination of photodissociation in the warm surface layers and freeze-out in the cold, dense parts of the disk. In these models, the molecular emission largely originates from the region just below the disk surface where the HCO^+ abundance climbs to a few times 10^{-10} . In general, the HCO^+ abundance follows the CO abundance in the disk, since its formation is primarily from CO reacting with H_3^+ , with destruction by dissociative recombination with free electrons.

Figure 5 shows the visibility amplitude of the $\text{HCO}^+ J = 1-0$ emission from TW Hya as a function of baseline length; the falloff at longer baselines demonstrates that the

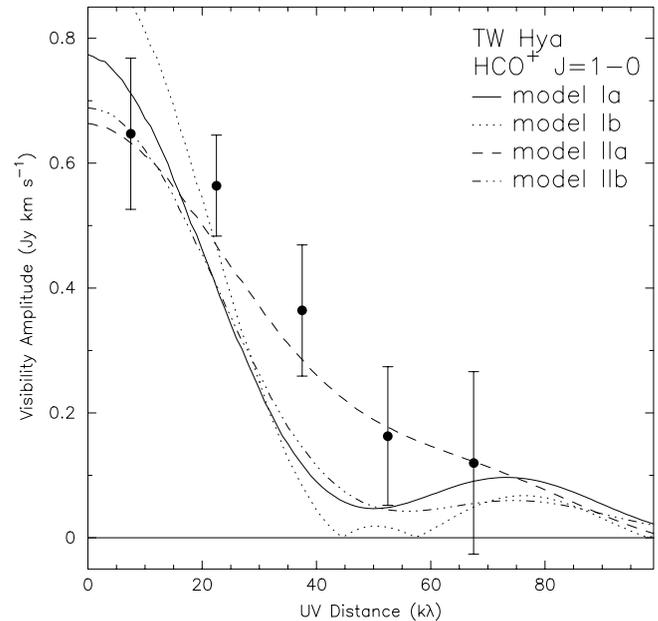


FIG. 5.—Visibility amplitude of observed and modeled $\text{HCO}^+ J = 1-0$ emission vs. baseline length, integrated over 0.5 km s $^{-1}$ width centered at LSR velocity 2.75 km s $^{-1}$, annularly averaged in 15 $k\lambda$ bins. The error bars represent ± 1 standard deviation for each bin. The four curves show the visibility amplitudes derived from the model calculations of van Zadelhoff et al. (2001) that fit the single-dish observations of higher J HCO^+ transitions. Models Ia and Ib are based on the radiatively heated disk models of D’Alessio et al. (1999), and models IIa and IIb are based on the passive two-layer models of Chiang & Goldreich (1997). For model Ia (solid line) and model IIa (dashed line), HCO^+ is depleted by an overall factor of 100, with an additional order of magnitude drop when the temperature falls below 20 K. For model Ib (dotted line) and model IIb (dash-dotted line), HCO^+ emission is depleted by a factor of 500.

emission region is resolved. A circular Gaussian FWHM “size” of $3''.2$ provides a crude description of the full line brightness distribution. Figure 5 includes the visibility amplitudes derived from four of the favored models of van Zadelhoff et al. (2001), which provide reasonable fits to higher J HCO^+ single-dish observations. Two models are based on the radiatively heated accretion disk structure of D’Alessio et al. (1999), and two models are based on the passive two-layer description of Chiang & Goldreich (1997). Two chemical scenarios are presented for each of these physical models. The interstellar HCO^+ abundance is assumed to be 5×10^{-9} , and in one model HCO^+ is depleted by a factor of 100, with an additional order-of-magnitude drop when the temperature falls below 20 K, while in another model HCO^+ is depleted by a factor of 500 throughout. Table 2 summarizes the models. Figure 5 shows that all of these models, which have disk masses of $0.03 M_\odot$ and disk radii of 200 AU, agree well with the ATCA observations, for both the size scale and the absolute intensity of the HCO^+ line emission. (Note that visibility amplitude is a positive definite quantity, which results in a positive bias in Fig. 5.)

Figure 6 provides another view of the four models in the form of synthetic images. These images were made by sampling each of the models with the same visibility distribution as obtained by the ATCA observations, taking a 0.5 km s $^{-1}$ bin centered on the line. These images can be compared with the panel in Figure 2 for LSR velocity 2.75 km s $^{-1}$, where

TABLE 2
DISK MODELS

Model	Physical Structure (mass $0.03 M_{\odot}$, radius 200 AU)	HCO ⁺ Depletion
Ia	Irradiated accretion disk ^a	100 × (1000 × for $T < 20$ K)
Ib	Irradiated accretion disk ^a	500 ×
IIa	Passive two-layer disk ^b	100 × (1000 × for $T < 20$ K)
IIb	Passive two-layer disk ^b	500 ×

NOTE.—See the discussion by van Zadelhoff et al. 2001 for more details of these models and their comparable fits to unresolved observations of higher J HCO⁺ transitions.

^a D’Alessio et al. 1999.

^b Chiang & Goldreich 1999.

the line emission peaks. In this view, it is not easy to see that the line emission is well resolved spatially, but the similarity of the models in overall brightness and spatial extent is clearly apparent.

While none of the models of van Zadelhoff et al. (2001) were fine-tuned to match all of the available constraints, and none provides a unique best fit to the ATCA data, the high HCO⁺ depletion factors are a robust feature. All four models are indistinguishable within the noise of the observations, although there are differences in detail in the visibility distributions among the models. There is perhaps a hint that the models that produce the more concentrated emission distributions are best. The model curves in Figure 5 suggest that resolved observations with better signal-to-noise ratios have the potential to discriminate between subtly different physical and chemical scenarios.

HCO⁺ depletions of an order of magnitude or more have been inferred for other T Tauri systems from millimeter imaging and analysis, in particular the disks around the GG Tau binary (Guilloteau, Dutrey, & Simon 1999) and around LkCa 15 (Duvert et al. 2000; Qi 2001). These isolated pre-main-sequence systems are located in holes of the Taurus dark cloud complex, which make their disks readily accessible without confusion. The radii of the gas disks in these systems, ~ 800 and ~ 600 AU, respectively, are among the largest known, and larger by factors of 3 to 4 than the ~ 200 AU radius disk surrounding TW Hya. That the depletion is similarly high in all of these disks provides further evidence that substantial depletion must occur in the dense shielded midplane material, rather than simply in the colder, outer parts of the disks.

Since the critical density of the HCO⁺ $J = 1-0$ line for collisional excitation is $\sim 6 \times 10^4 \text{ cm}^{-3}$, the detection of extended emission indicates high densities must be present to large radii, independent of any detailed physical and chemical model for the disk. This in situ measurement of local density confirms indications from scattered-light models and also analysis of dust emission that relied on extrapolation of the inner disk density distribution. Resolved images of multiple transitions of multiple species may be used to constrain the disk mass (see Dutrey, Guilloteau, & Guelin 1997).

4.2. HD 100546

For HD 100546, the spherical radiative transfer models of Henning et al. (1994) that reproduce the far-infrared

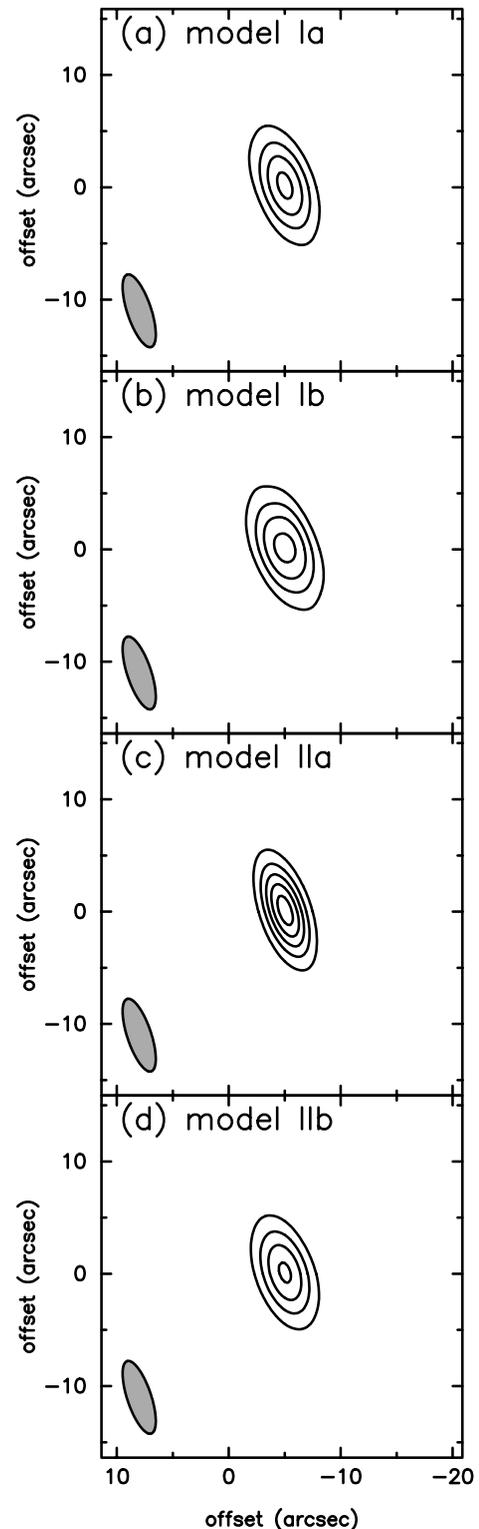


FIG. 6.—Synthetic images of the HCO⁺ $J = 1-0$ emission for the four physical and chemical models whose visibility amplitude distributions are shown in Fig. 5 (see § 4 for descriptions). For each image, the velocity width is 0.5 km s^{-1} . The contour levels are $\pm 2, 4, 6, \dots \times 66 \text{ mJy}$, as in Fig. 2. The ellipse in the lower left corner of each panel shows the synthesized beam.

emission do not account for the observed millimeter flux, even with a population of “fluffy” dust grains, which suggests that a circumstellar disk component must be present in the system. The compact 89 GHz dust continuum emission detected in the ATCA observations provides direct evidence

for this disk. An extrapolation of the 1.3 mm flux of 0.69 Jy (Henning et al. 1998) to 89 GHz gives 40 mJy, for a dust opacity power-law index near unity, compared to our detected flux of 36 ± 3 mJy. Thus, the disk component nominally accounts for a substantial fraction of the detected 89 GHz flux, although the uncertainties are large. Detailed modeling of the disk properties awaits additional millimeter and submillimeter observations with sufficient angular resolution to separate the disk and envelope components.

The similar continuum fluxes of HD 100546 and TW Hya suggest that the disk in the more distant and warmer HD 100546 system is substantially more massive than the disk around TW Hya. Since HCO⁺ $J = 1-0$ emission is not detected from HD 100546 with comparable sensitivity to the TW Hya detection, it is likely that the abundance reduction is even more extreme. Perhaps photodissociation of CO in the disk atmosphere is enhanced in the environment of this more massive, hotter star.

Grady et al. (1997) suggest that much of the accreting gas seen in the optical along the line of sight to HD 100546 is associated with star-grazing planetesimals, such as comets or asteroids. This may explain both the lack of observed boundary-layer emission, which might be expected from accretion of mostly gaseous material, and signatures of bipolar outflow in forbidden line emission that is characteristic of an optically thick, gaseous accretion disk. This fact, together with the nondetection of HCO⁺ emission and the highly processed nature of the dust around HD 100546, perhaps indicates that the gas is nearly fully depleted in the disk, in which case the planet-building process may be well advanced.

5. CONCLUSIONS

We have made high-resolution observations of the nearby young stars TW Hya and HD 100546 with the ATCA using three antennas equipped at 89 GHz. These observations

indicate the promise for future millimeter-wave observations of protoplanetary disks located deep in the southern sky. The observations of the TW Hya disk resolve emission from the HCO⁺ $J = 1-0$ line, and the measured extent and brightness support disk models with high depletion factors inferred from previous, spatially unresolved measurements of higher J transitions. The detailed modeling of the TW Hya disk represents a first step toward probing the gas content of this important system, whose dispersal and chemistry may be similar to the early solar system. A more complete understanding of the details of disk chemistry will require spatially resolved images with higher sensitivity, eventually sampling species that probe all of the relevant chemical families.

Thermal dust continuum emission from these sources was easily detected and promises to be resolved with observations using longer baselines. The available ATCA antenna stations will offer baselines up to 1.5 km in length, corresponding to an angular resolution better than $0''.5$ at this frequency, which will be adequate to probe the disk surface density and to start to disentangle disk structure from opacity effects. Both TW Hya and HD 100546 will be prime protoplanetary disk targets for the Atacama Large Millimeter Array, the next generation millimeter interferometer sited in northern Chile, construction of which has just started and will continue for the next decade.

We thank an anonymous referee for several suggestions that improved this paper. The Australia Telescope Compact Array is part of the Australia Telescope, which is funded by the Commonwealth of Australia for operation as a national facility managed by CSIRO. Partial support for this work was provided by NASA Origins of Solar Systems Program grant NAG5-11777, by the NRAO Foreign Telescope Travel Fund Program, by a NOVA network 2 grant, and by a NWO-Spinoza grant. C. M. W. acknowledges support of an Australian Research Council Fellowship.

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